

MRO/CRISM Retrieval of Surface Lambert Albedos for Multispectral Mapping of Mars With DISORT-Based Radiative Transfer Modeling: Phase 1—Using Historical Climatology for Temperatures, Aerosol Optical Depths, and Atmospheric Pressures

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Abstract—We discuss the DISORT-based radiative transfer pipeline (“CRISM_LambertAlb”) for atmospheric and thermal correction of MRO/CRISM data acquired in multispectral

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mapping mode (~ 200 m/pixel, 72 spectral channels). Currently, in this phase-one version of the system, we use aerosol optical depths, surface temperatures, and lower atmospheric temperatures, all from climatology derived from Mars Global Surveyor Thermal Emission Spectrometer (MGS-TES) data and from surface altimetry derived from MGS Mars Orbiter Laser Altimeter (MOLA). The DISORT-based model takes the dust and ice aerosol optical depths (scaled to the CRISM wavelength range), the surface pressures (computed from MOLA altimetry, MGS-TES lower atmospheric thermometry, and Viking-based pressure climatology), the surface temperatures, the reconstructed instrumental photometric angles, and the measured I/F spectrum as inputs, and then a Lambertian albedo spectrum is computed as the output. The Lambertian albedo spectrum is valuable geologically because it allows the mineralogical composition to be estimated. Here, I/F is defined as the ratio of the radiance measured by CRISM to the solar irradiance at Mars divided by π ; if there was no martian atmosphere, I/F divided by the cosine of the incidence angle would be equal to the Lambert albedo for a Lambertian surface. After discussing the capabilities and limitations of the pipeline software system, we demonstrate its application on several multispectral data cubes—particularly, the outer reaches of the northern ice cap of Mars, the Tyrrhena Terra area that is northeast of the Hellas basin, and an area near the landing site for the Phoenix mission in the northern plains. For the icy spectra near the northern polar cap, aerosols need to be included in order to properly correct for the CO_2 absorption in the H_2O ice bands at wavelengths near $2.0 \mu\text{m}$. In future phases of software development, we intend to use CRISM data directly in order to retrieve the spatiotemporal maps of aerosol optical depths, surface pressure, and surface temperature. This will allow a second level of refinement in the atmospheric and thermal correction of CRISM multispectral data.

Index Terms—Atmospheric propagation, infrared spectroscopy, remote sensing, software verification and validation.

I. INTRODUCTION

THE MARS Exploration Rovers, Spirit and Opportunity, have been blazing two different trails on the martian

surface for almost two Mars years now. Spirit has found that the Gusev Crater does not contain sedimentary deposits from Ma'adim Valles as initially believed, but rather that basaltic materials cover the crater floor [51] and the Columbia Hill outcrops exhibit a variety of interesting water-related alteration minerals [53]. Opportunity has confirmed the orbital-based detection of coarse-grained gray hematite at Terra Meridiani [52], and it has explored the depths of several small to medium-sized craters, allowing extensive stratigraphic analyses of the sediments [20], [52]. The instrumentation on both rovers has allowed truly spectacular mineralogical studies to be completed and the geological understanding of two sites on Mars to be markedly improved, thus giving ground truth for several orbital missions' studies of the martian atmosphere, surface, and subsurface (see, e.g., [5], [10], [46], and [47]).

These two rovers have blazed trails for future ground-based missions as well. The experience gained from these missions, both from an engineering perspective and from a scientific perspective, will offer (or has offered) greater probabilities of success for the Phoenix Lander 2007, for the upcoming Mars Science Laboratory (MSL) 2009, and for Exomars 2013.¹ However, from a geological and mineralogical perspective, the impact of the current and past orbital missions may do more to assist the upcoming landed missions than the spatially limited observations from either Spirit or Opportunity. From the bird's-eye perspective, orbital scientific instruments, such as HiRISE, HRSC, MGS-TES, THEMIS, MOLA, MOC, PFS, GRS, MARSIS, SHARAD, OMEGA, and CRISM, have the capability to characterize and evaluate future landing sites. These orbital instruments have global mapping capabilities of sufficient resolution to have been critical in the selection of a safe and scientifically interesting landing site for the ice-prospecting operations of the Phoenix Lander 2007 [49]. Similar geomorphic and mineralogical studies of data acquired with instrumentation in Mars orbit are currently providing a wealth of interesting landing site options for the MSL 2009.

The CRISM imaging spectrometer has been acquiring data from its platform on the Mars Reconnaissance Orbiter (MRO) since September 2006 [31]–[33]. The data consist of as follows: 1) hyperspectral targeted observations with high spatial resolution (~ 18 m/pixel) and 544 spectral channels, and 2) multispectral mapping strips with lower spatial resolution (~ 200 m/pixel) and 72 spectral channels. The spectral coverage for both modes of observation is $0.362\text{--}3.920\ \mu\text{m}$. The data from both observing modes have been extremely valuable to the Mars exploration community in the selection of landing sites for the MSL [17], [44], particularly due to the capability of CRISM to identify minerals (i.e., phyllosilicates, sulfates, etc.) on the surface of Mars. The high spatial resolution of CRISM is particularly useful for the landing site evaluation. It should be stressed, however, that CRISM has other capabilities besides landing site selection, including the comprehensive

spectral targeting and mapping of the entire planet in the near-infrared at significantly higher spatial resolution than before. This high spatial resolution in the near-infrared will allow for the detection of phyllosilicates, sulfates, and other minerals which contain H_2O ,² among other minerals.

In order to quantitatively understand the data from any instrument, investigators need to account for variable observing conditions [1], [2], [19], [26], [37], [38], [41]. From Mars orbit, an instrument that acquires data at visible and near-infrared wavelengths encounters variable photometric angles (solar incidence angle and viewing angle), atmospheric conditions (surface pressure and dust and ice aerosol optical depths), and thermal conditions (surface temperature and surface slope). In this paper, we will describe the photometric, atmospheric, and thermal correction system, as first discussed by McGuire *et al.* [30], that is being used for converting the calibrated CRISM multispectral mapping I/F data [27], [31]–[33] to a true representation of the surface reflectance spectra. This basic pipeline may provide a useful recipe for similar correction pipelines for visible-to-infrared instrumentation that observes Mars, either from a rover, from Mars orbit, or from the Earth/Moon system. Furthermore, the visible-infrared instruments are common to many missions, and this pipeline could be generalized to other solar system targets. After accounting for variable observing conditions, mapping and data mining of both the atmosphere and surface of Mars (e.g., [4], [5], [7], [10], [11], [27], [36], [42], [46], and [47]) can begin in earnest, without concern for the distorting effects of changing surface geometries and atmosphere.³ With proper correction, we can overlay or mosaic CRISM mapping image strips from different orbits at different times of the martian year of nearby locations, and the resulting digital map will have much reduced discontinuities between the different image strips. Such an endeavor will enhance our ability to accurately map [27] the surface properties of Mars and allow a more robust evaluation of future landing sites.

The I/F data from CRISM consist of the measured radiance-on-sensor I divided by the solar radiance F which is the solar irradiance (J) per steradian: ($F = J/\pi$). The I/F data from CRISM are measured simultaneously for up to 544 different wavelengths from 0.362 to $3.92\ \mu\text{m}$, giving a spectrum in the visible to near-infrared which allows studies of both the minerals in the surface and the constituents in the martian atmosphere [31]–[33]. After the photometric, atmospheric, and thermal correction of CRISM multispectral I/F data, the resulting A_L data are the incidence-angle-corrected reflectance called “Lambertian albedo.”⁴ The retrieval of spectra of Lambert albedos allows the following: 1) the estimation of mineralogical

² H_2O ice has strong absorptions at 1.5 and $2.0\ \mu\text{m}$. When the mineral is not ice but, instead, contains H_2O either in a chemically bound or physically bound form (i.e., hydrated phyllosilicates and sulfates), then these absorptions at 1.5 and $2.0\ \mu\text{m}$ are modulated and shifted to different wavelengths in the near infrared.

³The variable atmospheric conditions, which are corrected in this paper, include varying atmospheric CO_2 pressure and varying dust and ice aerosol optical depths. By correcting these major effects of the atmosphere, we can more readily assess more subtle atmospheric components of water vapor, carbon monoxide, oxygen/ozone, etc.

⁴Hereafter, we will often use the term “Lambert albedo” instead of “Lambertian albedo,” for short.

¹According to recent press reports in October 2008 (i.e., http://news.yahoo.com/s/space/20081017/sc_space/europetodelaymarsrovermission), “European Space Agency (ESA) governments have tentatively agreed to delay the launch of the Exomars rover to 2016, as part of a broader effort to rein in project costs and seek deeper cooperation with NASA and the Russian space agency.”

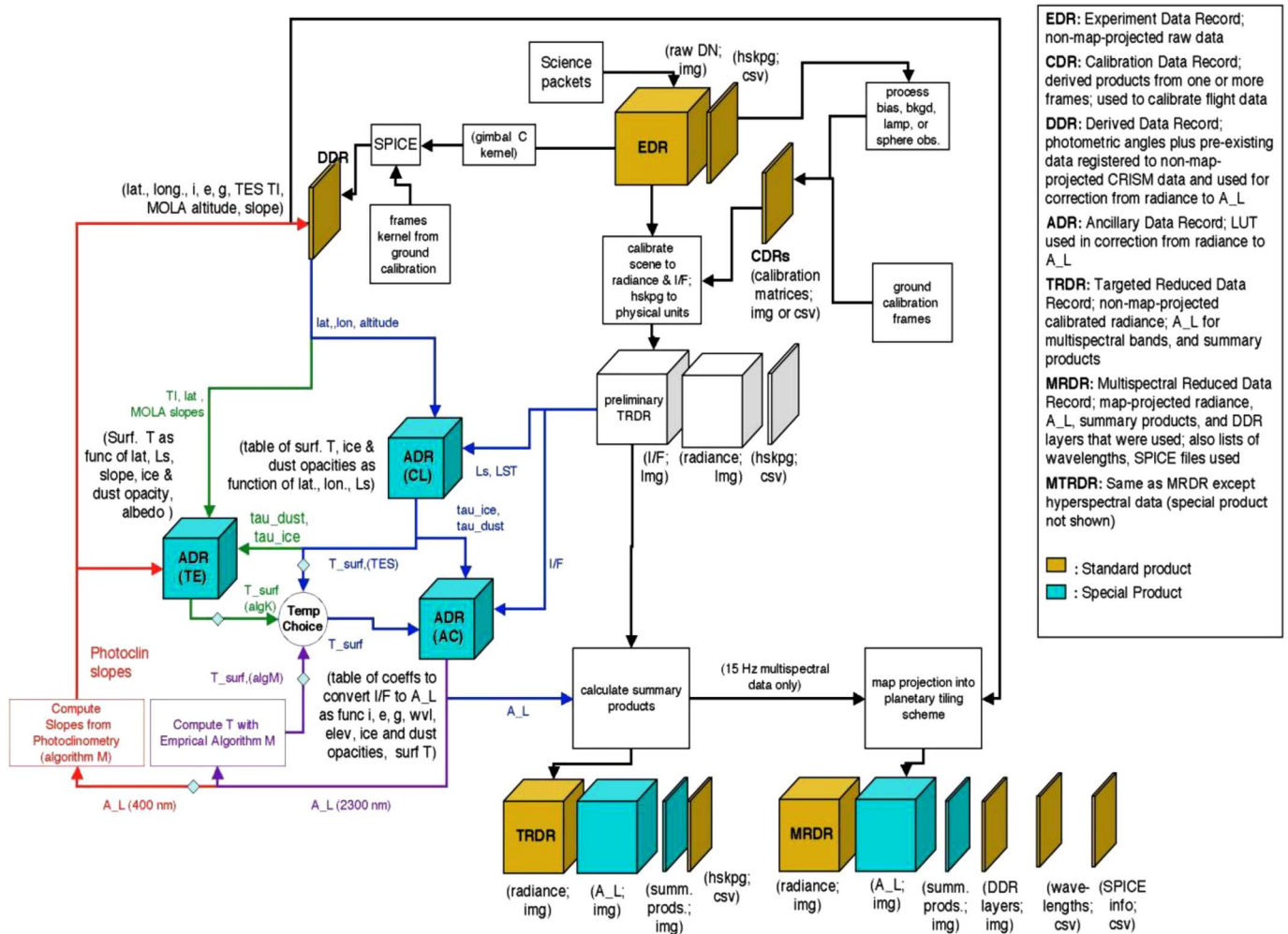


Fig. 1. Software flowchart [31] for converting the raw data (in the EDR) to calibrated I/F data (in the preliminary TRDR) to the photometrically, atmospherically, and thermally corrected Lambert albedo data (in the A_L TRDR). The DDR is used for backplane information such as the photometric angles and the georeferencing latitudes and longitudes. The ADRs are used as LUTs to query during the atmospheric and thermal correction; these ADR LUTs connect A_L to radiance (and to I/F), i.e., the LUTs connect actual surface properties to spacecraft measurements. The MRDR⁵ is computed by map-projecting a number of Lambert albedo TRDRs into a given map tile. In phase one of this project, we use climatological thermometry of the surface from the MGS-TES spectrometer. Therefore, the complexity of the thermal correction shown in this flowchart is reduced considerably. In this paper, therefore, we focus particularly on the blue lines.

abundances and 2) the quantitative intercomparison of spectra from different orbits of the spacecraft under different observing conditions. The photometric model for a flat Lambertian surface without an atmosphere is $I/F = A_L \cos(\text{INC})$, where INC is the solar incidence angle that is relative to the surface normal [22, pp. 190–191]. This photometric model presumes that the diffusely reflected photons have equally probable surface-reflectance angles, called the emission angles⁶ (EMI). This independence of I/F on EMI often occurs for bright granular surfaces. Other more sophisticated non-Lambertian photometric models have been developed, for example, the Lunar-Lambert model [29] or the Hapke model [22], but for our multispectral mapping work with CRISM, we will begin with the simpler Lambertian model, with the intent that we will later extend this paper to non-Lambertian models, as warranted.

For a surface with an atmosphere overlying it, the correction is somewhat more complex, particularly for variable observing conditions. The effects of atmospheric scattering and absorption by both molecules and aerosols over different surfaces have been analyzed previously (for example, see the works of Bohren and Clothiaux [6], Chandrasekhar [9], and Thomas and Stamnes [55]).

Further detailed work on radiative transfer modeling through the martian atmosphere and off of surfaces of constant albedo has been done in [2, Fig. 1] for a case where the optical depth of dust aerosols is 0.35 and the optical depth for ice aerosols is 0.05, with the aerosol optical depths measured at the MGS-TES reference wavelengths of 9.3 and 12.1 μm , respectively. The modeling work by Arvidson *et al.* [2] was done in order to quantitatively extract Lambert albedos from data from the OMEGA spectrometer. Arvidson *et al.* [2] found that for these aerosol optical depths, for “darker” surfaces of Lambert albedo less than a “transition Lambert albedo” of $A_{L,\text{trans},1} \sim 0.15$, the modeled I/F is higher than the Lambert albedo by

⁵Map-projected Reduced Data Record.

⁶Also known as “emergence angles.”

up to 80%, with the greatest effect for the lowest albedos.⁷ Arvidson *et al.* [2] also found that for these aerosol optical depths, for “brighter” surfaces of Lambert albedo greater than the transition Lambert albedo of $A_{L,trans,1} \sim 0.15$, the modeled I/F is lower than the Lambert albedo by up to 20%, with the greatest effect for the highest albedos. These changes in the I/F from the expected values without an atmosphere are caused by the aerosols. These general statements apply for wavelengths between 0.6 and 2.6 μm , exempting the CO_2 gas bands between 1.9 and 2.2 μm . For shorter wavelengths between 0.4 and 0.6 μm , a similar behavior applies but with different amplitudes and a transition Lambert albedo of $A_{L,trans,0} = 0.08$ instead of the longer wavelength transition Lambert albedo of $A_{L,trans,1} = 0.15$. For wavelengths in the CO_2 gas bands between 1.9 and 2.2 μm , Arvidson *et al.* [2] modeled that the absorption can be rather deep, with a 50% relative absorption by CO_2 for surfaces of Lambert albedo equal to 0.35 and with a 30% relative absorption by CO_2 for surfaces of Lambert albedo equal to 0.05.

All in all, this variance of behaviors of relatively darker I/F for bright surfaces and relatively brighter I/F for dark surfaces (in the presence of aerosols) is a primary reason for the detailed radiative transfer correction of the CRISM data that we present here. Aerosols reduce contrast; the radiative transfer correction of CRISM data restores much of this contrast. Furthermore, the amplitude of darkening over “brighter” surfaces and the brightening over “darker” surfaces depends on the aerosol content in the martian atmosphere as well as the photometric viewing angles [2], [59]. Wolff *et al.* [59] applied a one-band version of the technique, which we describe here, for MGS-TES broadband albedos over the nonpolar regions of Mars (60° S–60° N), and found that the bright areas of Mars have solarband MGS-TES Lambert albedos that are typically 5% brighter than without atmospheric correction (see [59, Fig. A1]). Wolff *et al.* [59] also found that dark areas of Mars have solarband MGS-TES Lambert albedos that are up to 30% darker than without atmospheric correction. Without radiative transfer correction of aerosol effects over surfaces of variable brightness, the quantitative accuracy of CRISM multispectral products would be significantly diminished.

The structure of this paper is as follows. First, we give a general overview of the pipeline processing used for CRISM multispectral mapping to correct for surface pressure, aerosols, thermal emission, and photometric angles. Second, we discuss the DISORT-based radiative transfer retrieval of Lambert albedos. Third, we describe the climatology of pressure, temperature, and aerosols, where “climatology” is defined as the historical record of these quantities at different places on Mars as a function of Mars year, solar longitude (day of year), and time of day. Fourth, we describe some results from the application of the correction system to several different CRISM multispectral image cubes or map tiles, including a particular

example of correcting a CRISM map tile for an area near the landing site for the Phoenix Lander 2007. Finally, we summarize our results and discuss future work.

II. GENERAL OVERVIEW OF PIPELINE PROCESSING USED TO CORRECT FOR SURFACE PRESSURE, AEROSOLS, THERMAL EMISSION, AND PHOTOMETRIC ANGLES

In Fig. 1, we show the processing flowchart for the photometric, atmospheric, and thermal correction system described in this paper. The multispectral data cube in raw EDR format⁸ is calibrated to I/F format (preliminary TRDR),⁹ which is then converted to a Lambert albedo (A_L) TRDR by correcting the following:

- 1) photometric effects of different viewing and incidence angles;
- 2) atmospheric effects of scattering and absorption by CO_2 and aerosols;
- 3) thermal effects caused by emission of photons by surfaces of different temperature.

The intent of this paper is to describe these different correction modules of our system and to give some examples from the initial applications of this system for retrieving Lambert albedos to multispectral CRISM data. In later papers, we will summarize more comprehensive results, system performance, and the utilization of this system for further scientific studies than those presented here. The scientific ends of this system include improvements in mapping the mineralogy of the martian surface and climatological mapping of gas abundances in the martian atmosphere [27].

In addition, in a later paper, we will discuss in more detail the system for temperature estimation that is summarized in Fig. 1. This paper focuses on the blue-colored lines in Fig. 1, which account for the correction for photometric angles, for surface pressure, and for aerosols. Climatological correction for thermal emission is also included as part of the current system, as outlined also by the blue-colored lines.

The purple-, red-, and green-colored lines in Fig. 1 represent how the correction for thermal emission using higher spatial-resolution techniques for temperature estimation will be phased into the CRISM_LambertAlb system. These higher spatial-resolution techniques for temperature estimation are not yet in use, largely since the thermal spectral bands (at wavelengths $> 3.0 \mu\text{m}$) have just recently been calibrated. Currently, we rely on the simplest choice of temperature, which is to use the ADR-CL¹⁰ lookup table (LUT) to look up the climatological value of MGS-TES surface temperature [47]. This is represented by a portion of the blue lines in Fig. 1.

We do not correct the MGS-TES surface temperatures for the ~ 1 -h difference in time-of-day between Mars Global Surveyor overflight and MRO overflight. These MGS-TES surface temperatures have $3^\circ \times 7.5^\circ$ spatial resolution in latitude and longitude [47]. The use of surface thermometry from MGS-TES is a proxy for more advanced surface thermometry, which will

⁷The term “transition Lambert albedo” means the value of the Lambert albedo (for particular levels of the ice and dust aerosols) that divides the brightened “darker” surfaces from the darkened “brighter” surfaces. $A_{L,trans,1}$ is defined as the value of this transition Lambert albedo for most wavelengths between 0.6 and 2.2 μm . $A_{L,trans,0}$ is defined as the value of this transition Lambert albedo for wavelengths between 0.4 and 0.6 μm .

⁸Experimental data record (EDR).

⁹Targeted reduced data record (TRDR).

¹⁰Ancillary data record—climatology.

be fully incorporated later into the Lambert albedo retrieval system for CRISM multispectral data. Enhancements to the spatial resolution of the surface temperature data will include as follows: 1) an empirical approach to estimate the albedos between 3.8 and 3.92 μm based upon an observed correlation between albedos at 2.5 μm and albedos at 3.8–3.92 μm ¹¹ (purple lines in Fig. 1), and 2) a physical approach to estimate the surface temperature for each pixel based upon using the ADR-TE¹² LUT which was computed from a physical thermal model of the martian surface (red and green lines in Fig. 1). The empirical thermal-correction approach [23] will allow the thermal emission to be estimated at 3.8–3.92 μm , which will facilitate the estimation of temperature for each CRISM spatial pixel, allowing thermal correction for shorter thermal wavelengths, i.e., between 3.0 and 3.8 μm . The physical thermal-correction approach [24], [28] has input parameters that include surface albedo, solar incidence angle, surface slope, surface slope azimuth, altitude, and aerosol optical depths. Both of these more advanced thermal-correction techniques will allow better identification of H₂O and CO₂ ices on the surface in minor quantities, for example.

III. DISORT-BASED RADIATIVE TRANSFER RETRIEVAL OF LAMBERT ALBEDOS

In order to estimate the Lambert albedo for a given CRISM spectel,¹³ we employ a LUT named the “Ancillary Data Record—Atmospheric Correction,” or ADR AC for short. This ADR-AC LUT was computed by using a DISORT-based¹⁴ radiative-transfer forward-model software package [54] to model how I/F will change as a function of 6–8 input parameters (see Fig. 2). For a given spectel at a given wavelength λ , the subroutine for accessing the ADR-AC LUT is called multiple times in order to “invert” the LUT and to estimate the Lambert albedo from I/F . These multiple different values of Lambert albedo are then used for populating the values of the grid-point corners of a multidimensional cube. Multidimensional interpolation is then employed to estimate the value of Lambert albedo for interior points that are not grid points. In this way, Lambert albedos can be calculated for all spectels in a multispectral map tile or image.

The use of these ADR-AC LUTs allows a substantial acceleration in the retrieval of the Lambert albedos than would be allowed by direct forward calls to DISORT, reducing the computation time per MSP¹⁵ TRDR strip (2700 rows, 60 columns, 72 spectral channels)¹⁶ on a fast Linux or

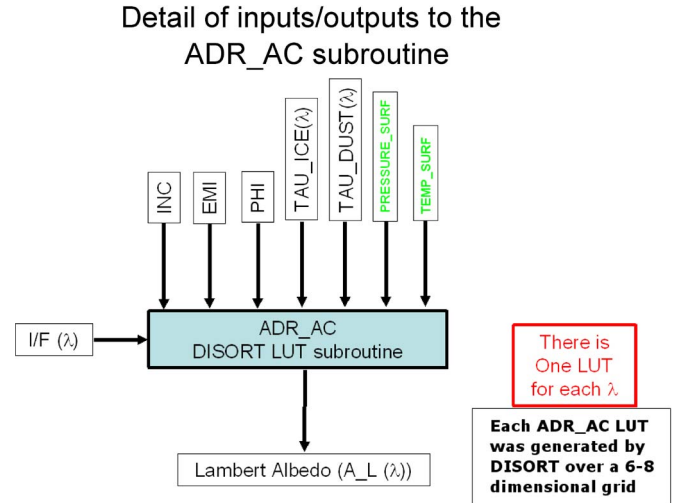


Fig. 2. Detail of how the ADR AC is queried. This LUT is shown in context in Fig. 1. The ADR AC consists of $I/F(\lambda)$ values computed with forward calls for each grid point at several input values of the photometric angles (INC, EMI, PHI), several input values of the aerosol opacities ($\tau_{\text{dust}}(\lambda)$, $\tau_{\text{ice}}(\lambda)$), and several input values of Lambert albedo $A_L(\lambda)$. The surface pressure and surface temperature inputs are only included as input values to the grid forward-call computation for particular wavelength bands. For use in the CRISM multispectral retrieval of Lambert albedos, the ADR-AC LUT is used in the inverse direction together with multidimensional interpolation, in order to use the $I/F(\lambda)$ value to estimate the Lambert albedo value $A_L(\lambda)$.

MSWindows workstation from about 3000 h¹⁷ to about 1 h. This is much slower than the calibration software that converts from the raw data to I/F . It is also much slower than a simpler alternative Lambert albedo retrieval software that corrects the I/F data for photometric effects by dividing by $\cos(\text{INC})$ and for atmospheric effects using the “Volcano-Scan” technique, without correcting for aerosol effects (for a discussion of some aspects of the Volcano-Scan technique, see Section V-B of this paper as well as the works of Langevin *et al.* [25] and Murchie *et al.* [32]). Nonetheless, the benefits of quantitative accuracy for this technique, which is based upon radiative transfer and utilizes a physical model for the correction of effects due to aerosols and for absorption in the CO₂ gas bands, can be worth the extra computation time when compared to the Volcano-Scan technique.

After the application of the atmospheric correction software to each TRDR, the TRDRs are mosaicked into map tiles (Fig. 1). At the equator, the tiles are on the order of 300 km \times 300 km in size, corresponding to $5^\circ \times 5^\circ$ in latitude and longitude. There will sometimes be more than 30 TRDR strips per map tile. With 1964 map tiles to span the planet and with about 61% CRISM coverage of the planet¹⁸ to date (as of March 1, 2008, with complete coverage expected), we need to make this Lambert albedo retrieval software as fast as possible, without sacrificing accuracy. Table I shows details on the grid parameters, whereas Table II provides detail of the atmospheric layers

¹¹Therefore, this empirical approach will need to wait until the CRISM data calibration is improved for wavelengths $> 3.7 \mu\text{m}$.

¹²Ancillary data record—thermal emission.

¹³We use the definition of spectel as the contents of a spectral channel for a particular pixel. For example, I/F spectel #1 is the value of I/F for spectral channel #1; spectral channel #1 is the channel of the L detector of CRISM which is at $\lambda = 3.92 \mu\text{m}$.

¹⁴DIScrete-Ordinates-method Radiative Transfer.

¹⁵Multispectral.

¹⁶This image size roughly corresponds to 548 km \times 13 km of surface area on Mars. The image width depends greatly on range to target. For the northern plains, the image width is about 13 km, but the MSP strips in the southern hemisphere are much closer to 10-km wide.

¹⁷By using 1 s for computing the Lambert albedo via forward calls for each spectel, this gives $(2700 \times 60 \times 72 \text{ spectels}) \times (1 \text{ s/spectel}) / (3600 \text{ s/h}) = 3240 \text{ h}$.

¹⁸The coverage depends on latitude: For the equatorial region, the coverage is 47%; the poles are nearly complete. Furthermore, this coverage does not account for periods of low data quality, most notably during periods of high dust optical depths in the atmosphere.

TABLE I

DEFINITION OF THE GRID POINTS USED BY DISORT TO GENERATE THE ADR-AC LUTS. FOR EACH OF THE GRID POINTS, DISORT IS RUN IN ORDER TO ESTIMATE I/F . THE CRISM_LambertAlb SYSTEM LATER USES THE ADR-AC LUT FOR EACH WAVELENGTH BAND TOGETHER WITH MULTIDIMENSIONAL INTERPOLATION TO ESTIMATE THE LAMBERT ALBEDO (A_L) FROM THE MEASURED I/F VALUE AND THE ESTIMATED VALUES OF THE OTHER DATA AXES (THE PHOTOMETRIC ANGLES: EMI, PHI, INC; THE AEROSOL OPACITIES: τ_{dust} , τ_{ice} ; AND THE SURFACE PRESSURE, THE SURFACE TEMPERATURE, AND THE SOLAR IRRADIANCE). THREE OF THE DATA AXES FOR THE GRID DEFINITION ARE ONLY USED WHEN EITHER THE SPECTRAL BAND IS A CO₂ GAS BAND OR A THERMAL BAND

Variable	# of grid points for data axis	Min. value	Max. value	Spacing
<i>The six grid axes that are common to all CRISM multispectral bands</i>				
Cos(EMI)	7	0.10	1.00	0.15
PHI (deg) (azimuthal angle)	11	0.00	180.00	18.00
Cos(INC)	7	0.10	1.00	0.15
A_L	11	0.00	0.60	0.06
$\tau_{\text{dust}}(9.3 \mu\text{m})$	9	0.01	0.71	0.0875
$\tau_{\text{ice}}(12.1 \mu\text{m})$	7	0.00	0.50	0.0833
<i>Extra DISORT grid axis used only for CO₂ gas bands: These gas bands are at: 1.21, 1.43, 1.66, 1.88, 1.97-2.17, 2.60, 2.70, and 3.00 μm</i>				
Surf Press.(mbar)	3	1.00	8.00	3.50
<i>Extra DISORT grid axis used only for thermal bands (3.12-3.92 μm):</i>				
Surf. Temp.(K)	3	180.00	300.00	60
<i>Extra DISORT grid axis used only for thermal bands (3.12-3.92 μm): The limits for two example thermal bands are shown (3.12 μm and 3.92 μm).</i>				
Solar Irradiance in $\text{W}/\text{m}^2/\mu\text{m}$ @3.12 μm	3	0.1113	0.1855	0.0371
Solar Irrad. in $\text{W}/\text{m}^2/\mu\text{m}$ @3.92 μm	3	0.0526	0.0877	0.0175

TABLE II

VITAL NUMBERS FOR THE ATMOSPHERIC MODEL USED BY DISORT IN ORDER TO COMPUTE THE ADR-AC LUTS. FOR THE ATMOSPHERIC EQUATION OF STATE, WE USE A REPRESENTATIVE TEMPERATURE PROFILE FROM THE MGS-TES DATABASE AND SIMPLY INTEGRATE THE HYDROSTATIC STRUCTURE EQUATION. GIVEN THE NEGLIGIBLE CONTRIBUTION FROM ATMOSPHERIC THERMAL EMISSION AND THE MINIMAL SENSITIVITY OF THE MOLECULAR ABSORPTION TO THE EXPECTED EXCURSIONS IN ATMOSPHERIC TEMPERATURES, OUR USE OF A REPRESENTATIVE PROFILE IS QUITE ADEQUATE. WE ALSO USED THE LAMBERTIAN PHASE FUNCTION (i.e., INSTEAD OF A HAPKE PHASE FUNCTION) FOR THE SURFACE PHOTOMETRY. FOR THE CO₂ GAS BANDS, WE ALSO EMPLOY THE CORRELATED-k TECHNIQUE (FOR A GENERAL DESCRIPTION, SEE [55])

Number of DISORT moments	Number of DISORT streams	Number of atmospheric layers
64	32	10
<i>Heights of boundaries between layers (km)</i>		
50.0	40.0	30.0
25.0	20.0	15.0
10.0	5.0	2.5
1.0	0.0	

currently being used for the CRISM multispectral retrieval of Lambert albedos. The grid point spacing and atmospheric layer spacing can be adjusted in future deliveries in order to give more accuracy or higher speed, but these settings have thus far been adequate for both accuracy and speed.

In Fig. 3, we show a portion of the internal structure of the ADR-AC LUT#2007.

This LUT was computed for the deep atmospheric CO₂ gas band at 2.007 μm . Note the straight-line dependence on semilog plots. This indicates that if we use exponential interpolation in the surface-pressure variable, the nonlinear dependence of I/F upon surface pressure can be adequately sampled by only three grid points for the pressure data axis. The exponential pressure scale heights of I/F , which we denote as $p_{\text{sh};I/F}$, range from ~ 5 to ~ 12 mbars for different values of the dust and ice optical depths. Although we do not show the dependences

of I/F on the other 5–7 data axes here, the majority of the parameter spaces for the other axes are dominated by nearly linear dependences. Because these dependences are not strictly linear, and occasionally nonmonotonic, we sample most of the other 5–7 data axes with more than three grid points. The plots in Fig. 3 also demonstrate that differences in ice and dust aerosol optical depths do have quite an effect on the coefficients to the exponential dependence of I/F on surface pressure. The use of multidimensional ADR-AC LUTs¹⁹ allows much faster computation than forward calls, and more accuracy than limited analytical or empirical models. This combination of speed and

¹⁹The ADR AC LUTs have sizes: 13.1 MB for each of the thermal bands, 4.4 MB for each of the CO₂ gas bands, and 1.5 MB for each of the normal bands. The sizes of these LUTs have been minimized to the extent possible, in order to optimize both the speed and the usage of memory.

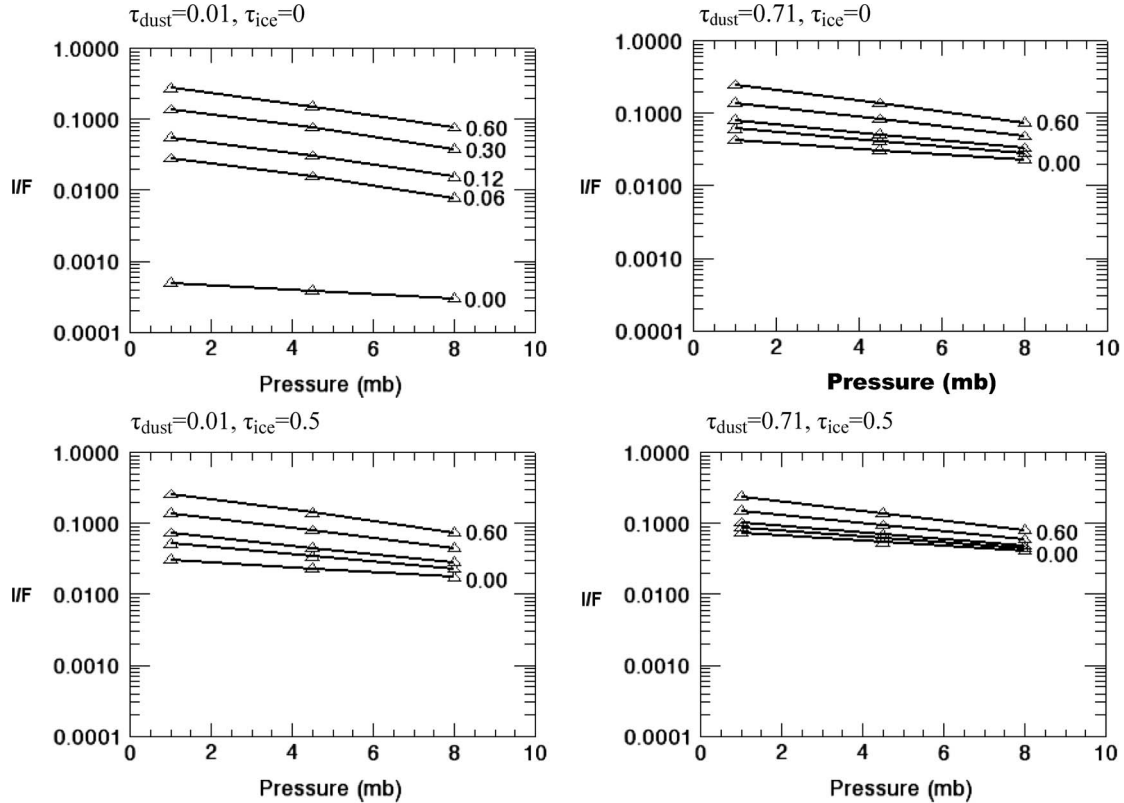


Fig. 3. We plot the dependences of I/F on surface pressure in semilog plots, in order to show a portion of the internal structure of the ADR-AC LUT, for $\lambda = 2.007 \mu\text{m}$. These dependences are given for five different values of Lambert albedo ($A_L = 0.0, 0.06, 0.12, 0.30, 0.60$) and four different sets of dust and ice aerosol optical-depth values. The indicated aerosol optical depths are for the MGS-TES reference wavelength. The straight-line semilog dependences suggest an exponential dependence, as is confirmed by detailed forward calls to DISORT (not shown here). The wavelength $\lambda = 2.007 \mu\text{m}$ is deep in the CO_2 gas band. These dust and ice optical depths are the values at $2.007 \mu\text{m}$, not the values for dust and ice optical depth at the mid-infrared reference wavelengths. For the input data axis of surface pressure, we find that three grid points form adequate sampling intervals from 1 to 8 mbars of surface pressure, but only if we use exponential interpolation for this data axis.

accuracy provides sufficient flexibility and robustness for the creation of the set of global map tiles of Lambert albedo for Mars.

As we will discuss in the next section on climatology, we use the estimates for the dust and ice aerosol optical depths from the TES instrument on MGS as inputs to the DISORT-based ADR-AC LUT in order to estimate the Lambert albedo spectra from the I/F spectra. However, the MGS-TES instrument measures the optical depths of the dust and ice aerosols at mid-infrared wavelengths: at the wavelengths of $9.3 \mu\text{m}$ for dust aerosols and $12.1 \mu\text{m}$ for ice aerosols. We have employed a Mie-scattering approximation (assuming spherical aerosols) [58] in order to convert the extinction cross sections of the aerosols at MGS-TES mid-infrared wavelengths to the extinction cross sections at CRISM wavelengths ($0.362\text{--}3.92 \mu\text{m}$).²⁰ This normalized extinction cross section is used multiplicatively to convert the aerosol optical depths measured at TES wavelengths to the aerosol optical depths measured at CRISM wavelengths. In order to generate scattering properties, for ice aerosols, we use the formulation from Warren [60] with $R_{\text{eff}} =$

$2.0 \mu\text{m}$ and variance $= 0.1 \mu\text{m}$, and for dust aerosols, we use $R_{\text{eff}} = 1.7 \mu\text{m}$ and variance $= 0.4 \mu\text{m}$ (for further discussion, see the work of Clancy *et al.* [11]).

IV. PRESSURE, TEMPERATURE, AND AEROSOL CLIMATOLOGY

In order to serve as our initial inputs to the DISORT-based Lambert albedo retrieval system, we are using the climatological records of aerosol optical depths and temperature from MGS-TES [47], as well as the surface altitudes from laser altimetry by the MOLA instrument [45], which is also on MGS. The MGS-TES climatology is in a LUT called “Ancillary Data Record—Climatology” (ADR CL), and for our initial work, we are using the MGS-TES climatology from Mars years 24 and 26 (1999–2000, 2003–2005).²¹ The MOLA altimetry is given for each pixel in the TRDR in a separate derived data record (DDR).

The MGS-TES surface temperature is used as an input to query the DISORT-based ADR-AC LUT in order to retrieve

²⁰The extinction cross section is defined as being the sum of the absorption cross section and the scattering cross section. The cross section of a particle is roughly equivalent to the effective area (i.e., of scattering or absorption) that a particle or a molecule presents to an incoming photon.

²¹The use of Mars years 24 and 26 is part of the ADR_CL operating mode named “Year of Best Data Quality”. Mars year 24 has the best data quality for TES, but Mars year 24 starts at $L_s = 103^\circ$. The data prior to $L_s = 103^\circ$ are coming from Mars year 26, since Mars year 25 had a dust storm.

Lambert albedos for the thermal bands at wavelengths $> 3.0 \mu\text{m}$. The optical depths for dust and ice aerosols are used by DISORT-based ADR-AC LUT to retrieve Lambert albedos for all spectral bands, but most of the effect is at wavelengths $< 1.2 \mu\text{m}$. The MGS-TES temperature of the lower atmosphere is used together with the MOLA altimetry and a Viking-based climatological record of surface pressure [56], in order to estimate the surface pressure for each spatial pixel in the TRDR, for different locations on the planet at different times in the Mars year. The MGS-TES lower-atmospheric temperature is computed by averaging over the four lowest available MGS-TES levels (above the surface), which is equivalent to averaging over 1 atmospheric pressure scale height.

In Fig. 4, we show the MGS-TES climatological record of the surface pressure, the temperatures, the aerosols, and the water vapor, for two longitudes and three solar longitudes²² as a function of latitude. This figure demonstrates the climatological spatiotemporal variability that is encapsulated in the ADR-CL LUT. Such variability implies two things: 1) Using a simple analytical or empirical model of most of these climatological variables is not sufficient for our purposes of retrieving Lambert albedos (hence, the importance of the ADR-CL LUT), and 2) with such past variability, it is unlikely that the ADR-CL LUT will always be correct (hence, the need for direct measurements of the climatological variables from the CRISM data). A secondary intent of this figure is to demonstrate that there is missing climatological data in the ADR-CL LUT. The absence of data is indicated in part by the dashed horizontal lines in these traces, particularly near the poles. The data value for the missing data locations and times is chosen to be the value for the nearest available location and time. Often, these nearest replacement values are inadequate for the accurate retrieval of Lambert albedos. This is particularly relevant for the missing aerosol values (i.e., near the poles). The aerosol values near the poles can be quite high and somewhat variable [34], [46], [47], [57], [62], and if their values are inferred from locations further from the poles, highly inaccurate retrievals of Lambertian albedos will ensue. Hence, we map the regions of missing aerosol data or poor aerosol data quality in Fig. 5. This map shows, for example, that the CRISM north polar mapping campaign will likely require, for accuracy, the direct retrieval of aerosol abundances instead of the nearest MGS-TES aerosol optical depths.

In the wavelength range of $1.2\text{--}2.6 \mu\text{m}$, the surface pressure is the most crucial parameter for correcting the data. Moreover, there are subtleties in estimating the surface pressure climatologically [15], [47], [50], [56], hence the emphasis in the discussion in this section on the surface pressure. The aerosols are primarily active shortwards of $1.2 \mu\text{m}$, but they do have low spectral-frequency effects at longer wavelengths, as well as interaction effects with the CO_2 gas bands at $2 \mu\text{m}$. The aerosols are determined directly from the MGS-TES climatological LUT (ADR CL). Therefore, with the exception

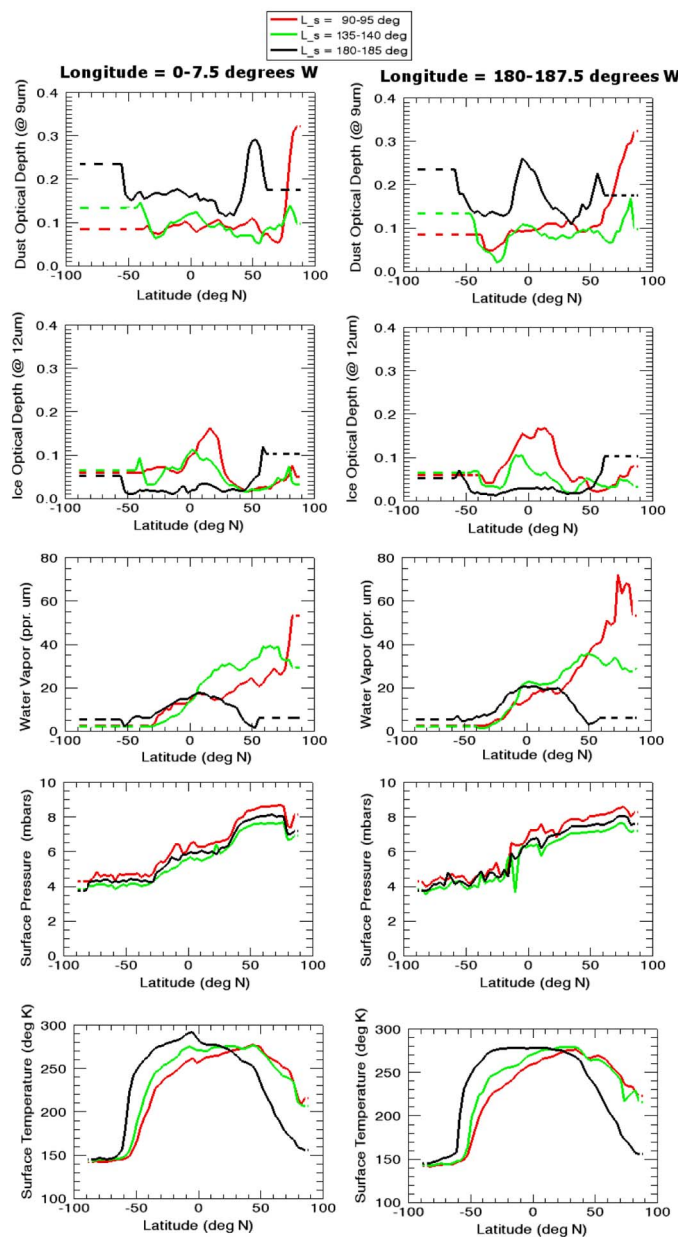


Fig. 4. ADR CL climatological profiles as a function of latitude for the indicated longitude ranges and the indicated solar longitude ranges. The dust and ice aerosol optical depths are shown at the mid-infrared MGS-TES reference wavelengths. Neither the ADR CL low spatial-resolution surface pressure nor the MGS-TES-estimated water vapor, both shown here, is used in the CRISM_LambertAlb system. We use the MGS-TES-estimated temperature of the lower atmosphere (not shown here), together with MOLA altimetry and Viking pressure climatology, in order to “algorithmically” estimate surface pressure at nearly the CRISM pixel scale. The MGS-TES-estimated surface temperature is used for the thermal correction for wavelengths $> 3.0 \mu\text{m}$. The dashed horizontal line segments represent gaps in the data (due to missing data or suboptimal data quality) that are currently filled in with nearby data.

of the rescaling, as described by Wolff and Clancy [58], and the data quality constraints shown later in Fig. 5, we limit the amount of discussion here for the aerosols, which is relative to the amount of discussion for surface pressure. Surface temperature for thermal correction longwards of $3 \mu\text{m}$ is also given directly by MGS-TES climatology in the ADR-CL LUT. Surface temperature has similar data quality constraints as the

²²Solar longitude (L_s) ranges from 0° to 360° and represents the position of Mars with respect to the Sun. The solar longitude of $L_s = 0^\circ$ corresponds to the spring equinox in the northern hemisphere on Mars.

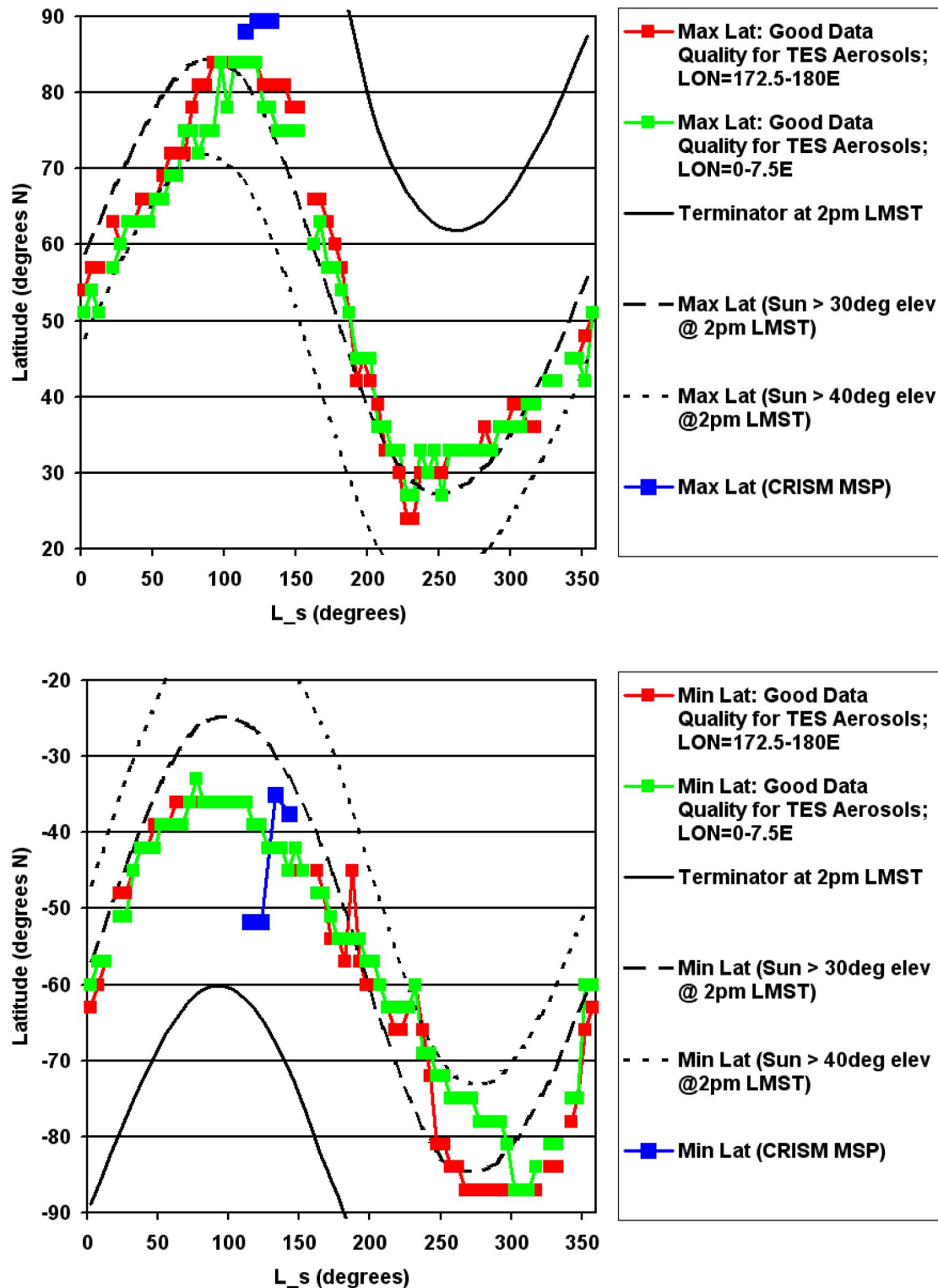


Fig. 5. We plot the latitudinal constraints of good MGS-TES-based aerosol climatology as a function of solar longitude (L_s) for two different planetocentric longitude ranges (0° – 7.5° E and 172.5° – 180° E) in order to show the northern most and southern most bounds for current phase *I* atmospheric correction of CRISM multispectral data. Temporal gaps in the aerosol climatology are shown by gaps in the green and the red curves. Also plotted are the positions of the terminator and the constraints on solar incidence angle as a function of solar longitude (L_s). The maximal extents of acquisition of CRISM multispectral data are also indicated, for the data that are currently available (August 2007) in the PDS geosciences data archive. Note that the maximal extents of MGS-TES aerosol data quality correspond roughly to the range of $> 30^\circ$ – 40° incidence angle for much of the martian year.

aerosols do. Therefore, we defer more detailed discussion of the aerosols and surface temperature inputs to the DISORT correction until a later phase in the CRISM multispectral

Lambert albedo retrieval project when we start retrieving aerosol optical depths and surface temperatures directly from the CRISM data.

TABLE III
MEAN PRESSURE

Mars year, Lander	Mean Pressure (mbar)
Year 1, VL1	7.936
Year 2, VL1	7.942
Year 1, VL2	8.663
Ave = $p(\text{VL})$	8.180

TABLE IV
AMPLITUDE OF FIRST FIVE HARMONIC TERMS OF PRESSURE CYCLE,
AVERAGED OVER THE FIRST TWO MARS YEARS FOR VIKING LANDER 1
AND THE FIRST MARS YEAR FOR VIKING LANDER 2

$a_1(\text{mbar})$	$a_2(\text{mbar})$	$a_3(\text{mbar})$	$a_4(\text{mbar})$	$a_5(\text{mbar})$
0.704	0.582	0.108	0.062	0.015

TABLE V
PHASE OF FIRST FIVE HARMONIC TERMS OF PRESSURE CYCLE,
AVERAGED OVER THE FIRST TWO MARS YEARS FOR VIKING LANDER 1
AND THE FIRST MARS YEAR FOR VIKING LANDER 2

	Φ_1	Φ_2	Φ_3	Φ_4	Φ_5
In Degrees	92.31°	-130.80°	-69.76°	-10.00°	49.55°
In radians	1.611 rad	-2.283 rad	-1.217 rad	-0.175 rad	0.865 rad

A. Algorithmic Pressure Climatology

The surface pressure on Mars varies seasonally (cf. [56]), and we must encapsulate that variation in an analytical “algorithmic” expression. This algorithmic climatology for surface pressure used by CRISM_LambertAlb is based upon an average of measurements of the two Viking landers over two Mars years without great dust storms (Viking Lander 1 for two Mars years and Viking Lander 2 for one Mars year, see [56]). The technique adopted by the CRISM team for averaging of the data from the two landers is detailed in the succeeding discussions and is similar to unpublished techniques used by the MGS-TES team. The pressure cycle over the course of a Mars year has characteristically two large and broad peaks, roughly preceding the winter solstice for each of the polar caps, after which CO_2 condensation on the caps begins in earnest. The Mars-year averaging is done for each of the five harmonic amplitudes and phases, and the averages are shown in Tables III, IV, and V. This pressure climatology as a function of solar longitude or Julian day is then offset to correspond to a MOLA altitude of 0 km. The “algorithmic” Viking-based climatological pressure equation used for CRISM multispectral retrieval of Lambert albedos is

$$p_{\text{surf}}(f, z = 0 \text{ km}) = p_{\text{surf,ave}}(z = 0 \text{ km}) \times (1 + (1/p(\text{VL})) \times \sum_k [a_k \sin(2\pi k f + \varphi_k)]). \quad (1)$$

In this equation, $p_{\text{surf,ave}}(z = 0 \text{ km})$ corresponds to the average surface pressure over a martian year for a surface that is at a MOLA altitude of $z = 0 \text{ km}$, and $p_{\text{surf}}(f, z = 0 \text{ km})$ is the time dependence of this surface pressure. The values of the amplitudes a_k and the phases φ_k are the averages in Tables IV

and V, the phases being given in radians. Seasonal time is expressed as f in terms of fraction of a martian year in seconds, not in L_s (i.e., f varies from 0 to 1). Note that L_s is not linear in time, so we use calendar date (or spacecraft clock time). In the aforementioned equation, f is equal to zero at $L_s = 330.2^\circ$, as in [56]. Relevant calendar dates for $L_s = 330.2^\circ$ include Nov. 26, 2005 (Julian Date 2453701) and Oct. 14, 2007 (Julian date 2454388).

The average for the three mean pressures is 8.180 mbar, but this is for the mean elevation of the Viking landers. The value of $p_{\text{surf,ave}}(z = 0 \text{ km}) = 5.477 \text{ mbar}$ that is in our formulation is determined by adjusting the $p(\text{VL}) = 8.180 \text{ mbar}$ mean value for the elevation of the Viking landers.

The pressure at other altitudes is computed from an exponential dependence with a given scale height H

$$p_{\text{surf}}(f, z) = p_{\text{surf}}(f, z = 0 \text{ km}) \exp(-z/H) \quad (2)$$

where the scale height depends on the temperature of the lower atmosphere: $H = RT/g$.

The gas constant is R , the gravitational acceleration is g , and the lower atmospheric temperature is T (in kelvin). For Mars, $H = (T/19.5) \text{ km}$, which for most daytime nonpolar conditions is roughly 10 km. The temperature of the lower atmosphere for a given location and time is determined from the MGS-TES climatology described earlier.

Using (1), we have calibrated for atmospheric conditions at the locations of the Viking landers 25–30 Earth years ago. This should allow somewhat accurate pressure determination at most locations on the planet, if we only know the altitude and the temperature of the lower atmosphere. This goes beyond what can be accomplished with the $3^\circ \times 7.5^\circ$ (in latitude and longitude) spatial resolution and 5° (in L_s) temporal resolution that we have made available in the ADR-CL LUT. The surface pressures in the ADR-CL LUT are also based upon Viking climatology, but the ADR CL uses a different definition of scale height than the one used in the algorithmic climatology here and is sampled at a much lower spatial resolution than the MOLA grid. The pressure climatology in the ADR CL is a rather crude estimate since it uses a constant 10 km for the scale height, and the climatology values are simply averages over the entire climatology bin from whatever observations happen to fall in that bin (which can be from very different altitudes than the multispectral pixel of interest). The differences between the two climatologies are usually less than 20% and often better than this. The largest systematic differences between the ADR-CL LUT climatology and the algorithmic climatology are observed in the Hellas basin, where the differences are greater than 1.5 mbars. A comparison is currently being performed between (a) the algorithmic pressure climatology, which is based in a large part on data from the Viking Landers, and (b) the pressures retrieved directly from CRISM data; this will be a topic in a future paper.

By inspection of Fig. 3 with the exponential dependence of I/F on surface pressure of a certain I/F scale height $p_{\text{sh};I/F}$ and by some calculus, a +1.5-mbar error in pressure, for a $p_{\text{sh};I/F} = 7.5 \text{ mbar}$, will have a –20% effect on the measured I/F . Therefore, if we want 5% photometry in the gas

bands for this particular case of $p_{\text{sh}, I/F} = 7.5$ mbar, we need 0.38-mbar accuracy in the pressure climatology. For certain locations on the planet, we will not be able to achieve this accuracy, one example being as follows: “in very rough terrain,” where there is significant variation in elevation at spatial scales smaller than the MOLA footprint of ~ 300 m. However, the current accuracy in high spatial-resolution Viking-based algorithmic pressure climatology is adequate for many purposes and over much of the planet and will serve CRISM data-processing needs for the time being until we upgrade the Lambert albedo retrieval system to retrieve the surface pressures directly from the CRISM data.

V. APPLICATIONS OF THE CRISM_LAMBERTALB SYSTEM

We show here a small subset of the results and analyses that we have completed thus far. The intent of these analyses is to elucidate some of the capabilities of the CRISM_LambertAlb system first described in [30] for the correction of multispectral CRISM data. The results that we describe here are as follows: 1) comparison of Lambert albedo spectra from DISORT retrieval with library reflectance spectra of minerals; 2) analysis of corrected spectra from near the northern ice cap and comparison of these corrected spectra with the “Volcano-Scan” technique; 3) intercomparison of map tiles in the Tyrrenia Terra region northeast of the Hellas basin; and 4) demonstration of map-tile production for a region north of the Phoenix Lander 2007 landing site.

The CRISM instrument contains two detectors, one for wavelengths from 0.362 to 1.04 μm (named the S detector) and another for wavelengths from 1.02 to 3.92 μm (named the L detector). The I/F data cubes for all of the analyses described here use calibration “TRR2,” which is the current calibration as of September 2007. The version of the Volcano-Scan software is CAT 6.0, which uses $\cos(\text{INC})$ photometric correction, together with a transmission spectrum of the martian atmosphere derived from taking a ratio of a TRR1 version of the I/F spectrum at the summit to the I/F spectrum at the base of Olympus Mons. The version of the CRISM_LambertAlb system is 20070809, which is the latest version.

CRISM_LambertAlb does not correct for atmospheric water vapor or carbon monoxide. These corrections can be done afterward with special runs of DISORT (which we do not present here), and they are perhaps best done after the main pipeline processing because of the following.

- 1) Water vapor and carbon monoxide are relatively minor species in the atmosphere. Twenty precipitable micrometers of water vapor is a typical column depth of water, and 700 ppm of carbon monoxide is typical for the martian atmosphere [48].
- 2) The banddepths and bandwidths of water vapor and carbon monoxide are relatively small compared to those of carbon dioxide.
- 3) Water vapor is somewhat variable in abundance even on small spatial scales, so the amount of water vapor may need to be inferred directly from the CRISM data on a pixel-by-pixel basis.

A. Comparison of Lambert Albedo Spectra From DISORT Retrieval With Library Spectra of Minerals

The CRISM mapping strip MSP00002838_07 in the far northern latitudes near the northern ice cap has some nice spectral and mineralogical diversity, which allows the quick analysis of the correction capabilities of CRISM_LambertAlb. The mineralogical diversity includes H_2O ice, gypsum-bearing dunes, and dust-covered surfaces. This mineralogical diversity gives rise to spectral diversity that includes bright spectra, dark spectra, and spectra that are bright at some wavelengths and dark at others. In Fig. 6(a), three different spectra from a single multispectral strip (MSP00002838_07) in the northern latitudes are shown. This strip was chosen because it has both icy outcrops and gypsum-rich dunes [39] within it.

In Fig. 6(b) and (c), we compare two Lambert albedo spectra from this strip with the U.S. Geological Survey (USGS) reflectance mineral spectra [14] for gypsum and water ice. We have also multiplied the CRISM spectra in these two figures by an arbitrary constant in order to better compare with the library spectra; this can account for the effect of possible spectrally bland darkening agents on the surface [12].

The CRISM MSP Lambert albedo spectrum in Fig. 6(b) has a strong absorption at 1.9 μm and an absorption edge at 2.4 μm , which is a hallmark of polyhydrated sulfate minerals [32]. The twice-hydrated mineral gypsum compares somewhat well with this CRISM spectrum, including the absorptions at 1.9 and 2.25 μm and the drop-off at 2.4 μm [40]. However, the CRISM data lack an absorption at 1.8 μm , and the 1.45- μm band for gypsum corresponds to a band at slightly longer wavelengths in the CRISM data. This CRISM spectrum may not correspond to gypsum, but with good certainty, it corresponds to a sulfate-rich mineral.

The case is much stronger in Fig. 6(c) that the CRISM spectrum corresponds to an area that is rich in H_2O ice, with the absorptions at 1.3, 1.5, 2.0, and 2.5 μm [13], [18], [21], [35], [60].

B. Comparison of DISORT Correction and Volcano-Scan Correction

Such spectral diversity allows a decent measure of testing of the overall quality of the atmospheric correction capabilities of our system, particularly in the presence of obscuring aerosols in the martian atmosphere. In Figs. 7 and 8, we compare two different atmospheric-correction techniques, the DISORT-based climatological system described in this paper and the Volcano-Scan correction system [25]. The Volcano-Scan correction system relies on a measurement of the atmospheric transmission by the CRISM instrument, which is accomplished by taking the ratio of a nadir-looking I/F spectrum acquired at the summit of Olympus Mons to the I/F spectrum acquired at the base of Olympus Mons. The Volcano-Scan correction technique rescales this transmission spectrum for variable surface pressure at other locations of the planet. The Volcano-Scan technique does not account for aerosol optical depths that vary with time or location on the planet, nor does it correct I/F spectra well when there are ices present on

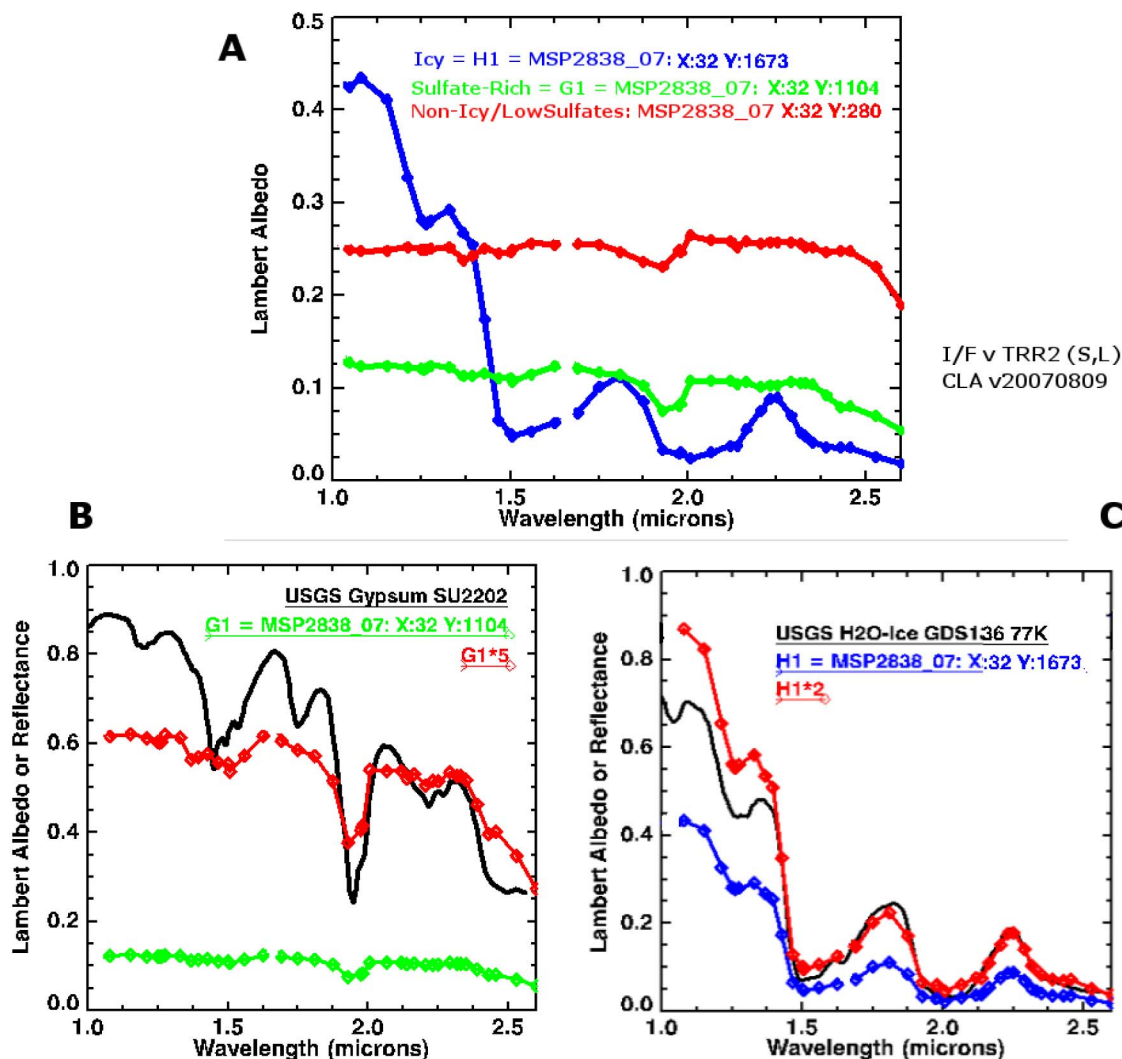


Fig. 6. In (a), three Lambert albedo spectra from CRISM multispectral strip MSP00002838_07. This strip is located near the north pole and crosses the gypsum dunes in Olympia Undae. These Lambert albedo spectra were computed by using the CRISM_LambertAlb software system. In (b), we show the sulfate-rich spectrum from (a), together with a Gypsum spectrum from the USGS library of reflectance spectra of minerals [14]. In (c), we show the ice-rich spectrum from (a), together with a H₂O-ice spectrum from the USGS library of reflectance spectra of minerals. Both (b) and (c) also contain versions of the CRISM spectra that are scaled by multiplying by an arbitrary constant, so that the CRISM spectra can be better compared with the USGS library reflectance spectra. The *X* and *Y* values in each subfigure are the locations of the pixel in CRISM multispectral image MSP00002838_07. For clarity, known channels of poor data quality (wavelengths of 1.02, 1.05, and 1.65 μm) have been removed from these spectra.

the surface. Fig. 7 shows this comparison between different atmospheric-correction techniques for bright spectra, and Fig. 8 shows this comparison for icy spectra that are both bright and dark.

Note that the atmospheric correction works well for both techniques in the CO₂ gas bands at 2.0 μm for the bright spectra. However, for the icy spectra which are dark in the broad H₂O ice bands at 1.5, 2.0, and greater than 2.4 μm , the DISORT-based method handles the CO₂ gas-band correction at 2.0 μm much more effectively. This latter deficiency of the Volcano-Scan technique (as currently implemented in CAT 6.0) is due to a combination of the unmodeled interaction of the aerosols at these dark wavelengths and the technique of fitting the banddepth at CO₂ when ices are present on the surface (H₂O and CO₂ ices also have deep bands at 2.0 μm). Also note the more subtle differences between the two atmospheric-

correction techniques over the entire spectral range from 1.0 to 2.5 μm in both figures. Again, this is due to the variable aerosols not being accounted for in the Volcano-Scan technique.

C. Comparison of Map Tiles for Tyrrhena Terra

A primary purpose of performing the correction of the CRISM multispectral *I/F* TRDRs for photometric, atmospheric, and thermal effects is to be able to mosaic or overlay multispectral TRDRs from different orbits under different observing conditions. If this correction was not performed prior to overlaying the different TRDR strips, then the varying observing conditions like different aerosol optical depths, different photometric angles, and different surface temperatures would cause the different TRDR strips, when juxtaposed, to appear wildly different. By performing this correction for different

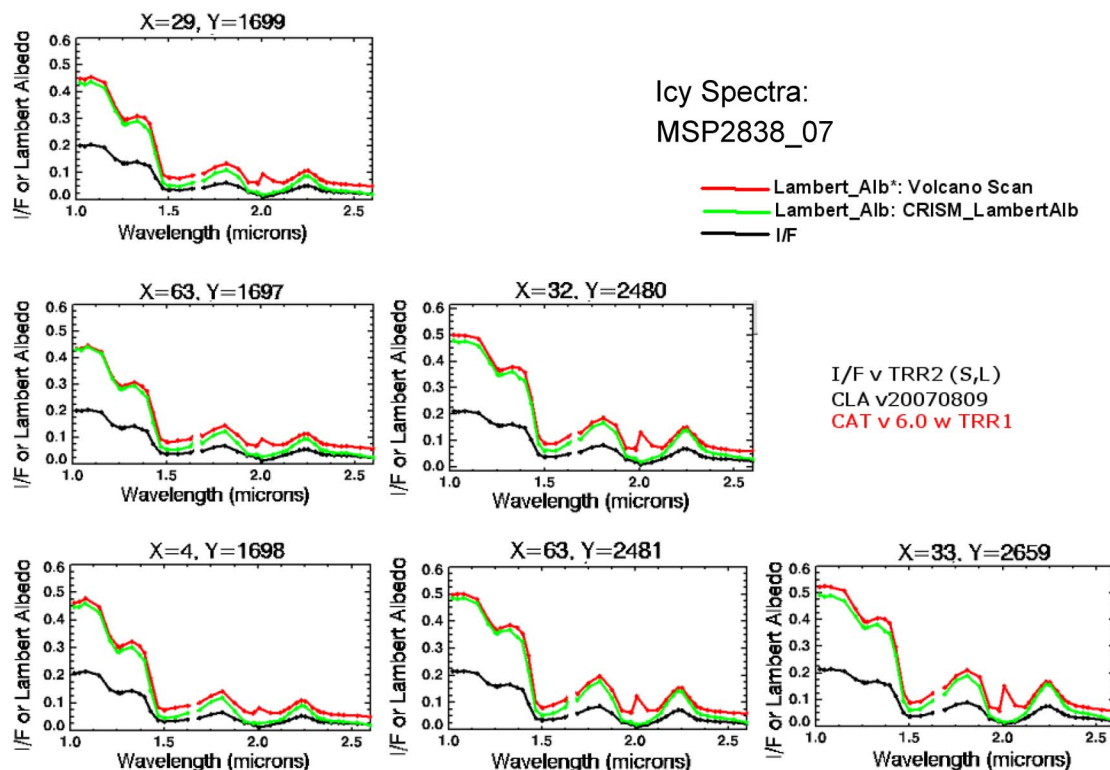


Fig. 8. We show some of the icy L-detector spectra from the CRISM Multispectral observation, MSP00002838_07, with the same parameters and description as in Fig. 7. For clarity, the instrumental artifact near $1.65\ \mu\text{m}$ has been removed from these spectra, although it is somewhat subdued in these icy spectra since the surface is darker at these wavelengths due to the broad H_2O -ice absorption band at $1.5\ \mu\text{m}$. The spectral channels near $2.0\ \mu\text{m}$ are well corrected by only the DISORT-based technique for these icy spectra, which are darker due to the deep H_2O -ice absorption band at $2.0\ \mu\text{m}$. The offset in Lambert albedo between the two correction techniques over the entire spectral range is caused by aerosol optical depths being explicitly handled by the DISORT-based technique and not handled by the Volcano-Scan technique. Note that the Lambert albedo spectra as computed by the DISORT-based technique (CRISM_LambertAlb) often appear to be closer to the I/F spectra for the darker wavelengths, but these same DISORT-corrected Lambert albedo spectra also often appear closer to the Volcano-Scan-corrected Lambert albedo spectra for the brighter wavelengths.

landing site. We show a portion of this map tile in Fig. 10, both without correction (I/F) and with correction (A_L). The quality of the correction is visible largely by visual inspection of the mosaicked TRDRs—the I/F version for the same image stretch has several TRDR strips that are dark and several that are bright, whereas for the A_L version, the contrast is much improved, showing at the right-hand side of the map tile some of the dark spots that are indicative of boulder fields surrounding craters, as first confirmed by the high-resolution HiRISE camera, also on MRO [3]. The histograms for the green channel (chosen here to be $1.506\ \mu\text{m}$) also show that the correction system converts the multimodal histogram²³ for the I/F map tile to a unimodal histogram for the A_L map tile, which is further evidence that the correction system is working well. The long small-amplitude tail for larger albedos than the main peak in the histogram of A_L map tile may be caused by slightly imperfect correction of some of the TRDR

strips. Seelos *et al.* [43] explore these map-tile mosaics in more detail.

Similar characterization of MSL landing sites (launch planned for 2009) with CRISM data (using both modes: multi-spectral mapping and hyperspectral targeted) is being pursued and already has been quite valuable in helping to narrow down the field of potential MSL sites to those that have interesting minerals, like phyllosilicates. Future landed missions like ExoMars (planned to be launched by the European Space Agency in 2013)²⁴ will also benefit greatly from mineralogical information afforded by CRISM, particularly with atmospheric and thermal correction, as outlined in this paper.

VI. FINAL REMARKS AND OUTLOOK

We have summarized in some detail the essential aspects of the system for correcting CRISM multispectral data for variable observing effects such as atmospheric aerosol optical depths, surface pressure, and surface temperature. The uncorrected CRISM I/F spectra are used along with the historical climatological observations of the aerosols, pressure, and temperature in order to retrieve Lambertian albedo spectra for each CRISM pixel. This system has been applied and tested on innumerable different CRISM multispectral data sets, and four of these tests

²³The multimodal histogram for the $1.506\text{-}\mu\text{m}$ channel of I/F is roughly a bimodal distribution with additional substructure for each of the main two peaks. The smaller of the main two peaks is caused by one MSP strip that has a roughly 7° lower incidence angle than most of the other strips in the map tile. This 7° lower incidence angle will make the I/F about 20% brighter for a Lambertian surface without an atmosphere. The observed difference in I/F between the lower incidence angle strip and an overlapping higher incidence angle strip is about 30%. Therefore, in this case, differences in incidence angle account for about two-thirds of the observed variation in I/F .

²⁴See footnote 1.

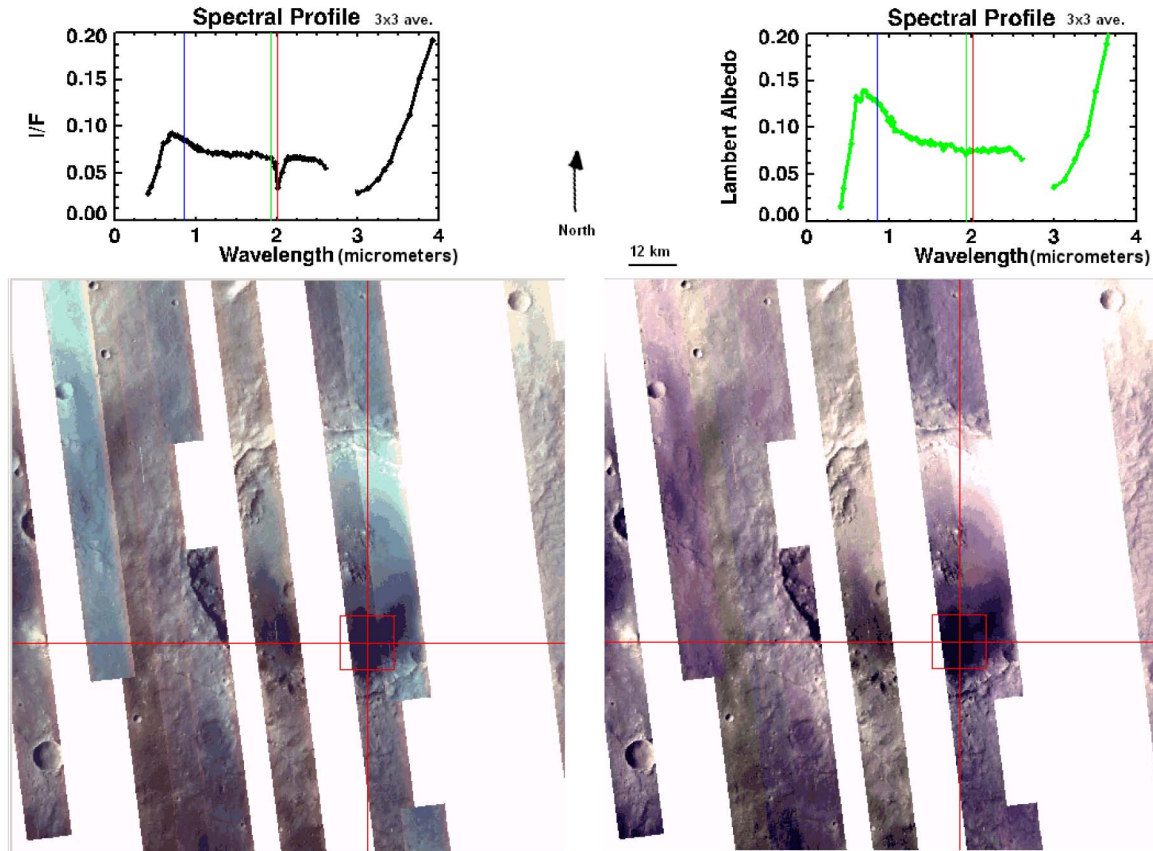


Fig. 9. CRISM multispectral map tile (with both the S and L detectors concatenated), T894, in the Tyrrhena Terra region, northeast of the Hellas basin, is shown, as corrected with the CRISM_LambertAlb DISORT-based software. On the left, we show a portion of the uncorrected I/F map tile and the spectrum. The lower left corner of this image is at 4.50° S, 96.88° E, and the upper right corner of this image is at 2.50° S, 98.69° E. On the right, we show the corrected version of the same map tile and spectrum, where we used standard pipeline processing and did not correct the photometric angles for surface slope. Thermal correction is not enabled here. North is upward in both images. Sun direction is somewhat variable, although a westerly incidence angle predominates. The bands shown here are $R = 2.01 \mu\text{m}$, $G = 1.92 \mu\text{m}$, and $B = 0.86 \mu\text{m}$. The stretching of the color space is min-max, with the bounding values chosen independently for each of the R , G , B planes for I/F and for Lambert albedo (I/F : R 0.032–0.086, G 0.066–0.155, and B 0.073–0.172; Lambert albedo: $R = G = B$: 0.10–0.30); the bounding values for the color stretching were not chosen to be the same for both I/F and Lambert albedo due to the change in range. The spectra shown are $3 \times 3 \text{ pixel}^2$ spatial averages centered at the point indicated by the cross-hairs in each map tile. Note that the blue coloring dominates for several of the strips in the I/F version, whereas this blue coloring is much reduced after atmospheric correction. Also note that the reddish coloring is common near the edge of most of the strips in the I/F version, whereas this red coloring is much subdued after atmospheric correction. This reddish coloring at the edge of the strips is caused by an optical distortion (spectral smile) in the CO_2 gas bands, for which we are explicitly correcting. For clarity, known channels of particularly poor data quality (wavelengths of $0.97\text{--}1.05$ and $2.66\text{--}2.80 \mu\text{m}$) have been removed from these spectra.

have been described here. Soon, we will be completing the production of the 1894 CRISM multispectral map tiles that span the planet.

In Phase II of this paper, which is currently being completed, we will add the physical thermal-correction approach [24], [28] discussed in Section II. This improvement will allow more accurate correction at smaller spatial scales at wavelengths greater than $3 \mu\text{m}$. Also, in Phase II, we will incorporate improved measurements of the aerosol scattering properties, as measured directly by CRISM [61]. This improvement will allow more accurate correction at wavelengths shorter than about $1 \mu\text{m}$.

After having applied the Phase I + II CRISM_LambertAlb system to the multitudinous multispectral TRDR strips, these map tiles will have much less additional unwanted variability due to shifting observing conditions. Hence, the true spectral diversity and variability can be identified more easily, documented, and understood. Indeed, the I/F versions of the map tiles are now being used in order to provide more targets for the CRISM hyperspectral targeting mode.

In the phases beyond Phase II, we will consider incorporating the following:

- 1) improved correction in the carbon dioxide gas bands, by directly sensing the carbon dioxide column depth for each pixel, following the techniques of the OMEGA team [15], [50];
- 2) water vapor correction and carbon monoxide correction, either as a part of the pipeline processing or as a stand-alone tool;
- 3) correction for non-Lambertian effects, for example, by using the Lunar-Lambert model [29] or the Hapke model [22];
- 4) aerosol correction of the polar regions, by incorporating maps of the polar aerosols, as measured by CRISM emission-phase-function measurements [8].

As the Lambert albedo versions of the map tiles are completed, spectral ambiguities will become disentangled, which

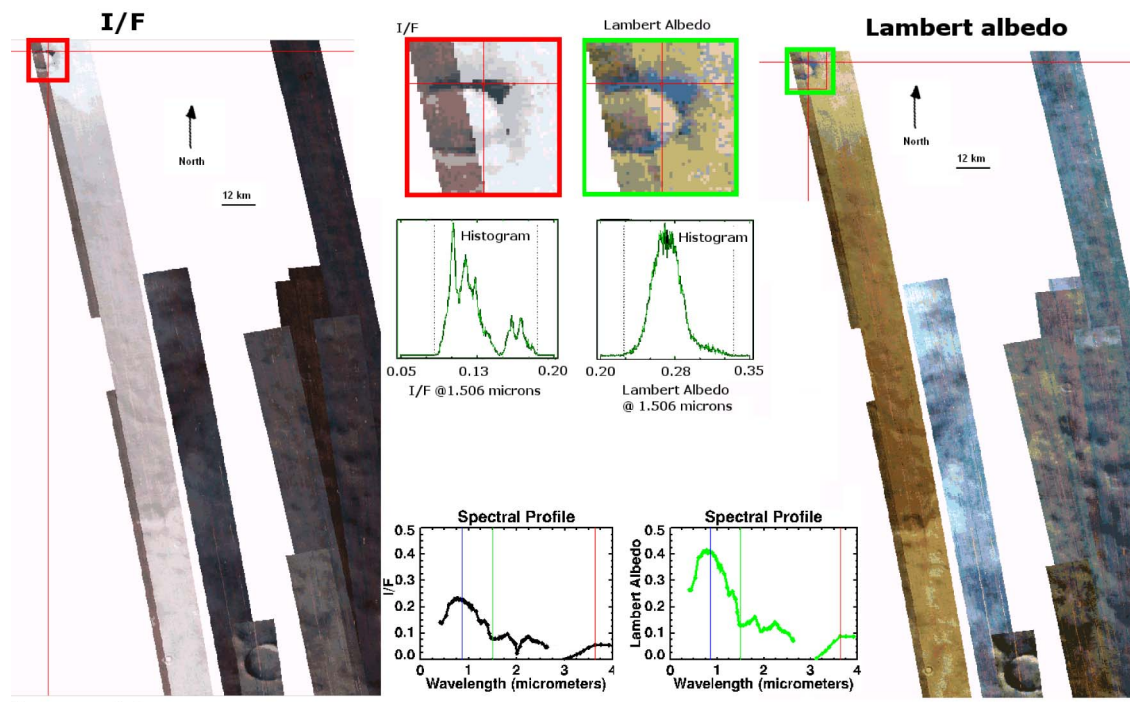


Fig. 10. CRISM multispectral map tile (with both the S and L detectors concatenated), T1889, near the Phoenix landing site region is shown, corrected with the CRISM_LambertAlb DISORT-based software. On the left, we show a portion of the uncorrected I/F map tile and the spectrum. The lower left corner of this image is at 69.21° S, 118.11° W, and the upper right corner of this image is at 72.50° S, 112.88° W. On the right, we show the corrected version of the same map tile and spectrum. These map tiles were constructed with a summer-only constraint on solar longitude (L_s between 105° and 165°). North is up in this image. The three color bands in the map tiles are $B = 0.83 \mu\text{m}$, $G = 1.506 \mu\text{m}$, and $R = 3.6 \mu\text{m}$. The stretching of the color space is min–max, with the bounding values chosen independently for each of the R , G , B planes for I/F and for Lambert albedo ($\{I/F\}$: R 0.067–0.173, G 0.083–0.184, and B 0.092–0.177; Lambert albedo: R 0.187–0.316, G 0.224–0.334, and B 0.240–0.364); the bounding values for the color stretching were not chosen to be the same for both I/F and Lambert albedo due to the change in range. Each strip in the tiles is approximately 12 km in width. In the middle of the figure, we show histograms over the whole map tile of the I/F and the Lambert albedo at $1.506 \mu\text{m}$. We also show zoom-ins of the ice-rimmed crater at 72.4° N, 117.5° W, north of the Phoenix landing site; this ice-rimmed crater is the location of the spectra shown here. There is no spatial averaging for these spectra. Note that the portion of the Lambert albedo map tile shown on the right has much better correction for photometric and atmospheric effects than the uncorrected I/F version on the left. This is apparent in several ways: 1) much better matching of the strips; 2) much less structure in the $1.506\text{-}\mu\text{m}$ histograms; and 3) much better visibility of some of the small dark rock fields in the right side of the Lambert Albedo map tile. For clarity, known channels of poor data quality (wavelengths of $0.97\text{--}1.05$, 1.65 , and $2.66\text{--}2.80 \mu\text{m}$) have been removed from these spectra.

will allow improved CRISM and HiRISE targeting of the martian surface during the primary science phase of the MRO. This disambiguation is particularly important near the CO_2 gas bands at a wavelength of $2.0 \mu\text{m}$, since there are a hydration band at $\sim 1.9 \mu\text{m}$ and a water ice band at $2.0 \mu\text{m}$. The atmospheric aerosol correction will be important for enabling mineralogical retrievals for the shorter wavelengths (shortward of $\sim 1.3 \mu\text{m}$), where pyroxene and olivine have different spectral characteristics. Furthermore, the thermal correction will allow the identification of different minerals that have absorptions between 3.0 and $3.9 \mu\text{m}$, including H_2O and CO_2 ices, carbonates, and organics. Carbonates and organics may not be particularly abundant at the multispectral mapping pixel scale of 200 m/pixel , but certainly the ices are abundant on these scales. The multispectral mapping resolution of CRISM will allow us to map these ices and other minerals at unprecedented spatial scales in the visible to near-infrared spectral range ($0.362\text{--}3.92 \mu\text{m}$). Hence, the next landed missions to Mars will have ever-increasing prospects for going to the right place for more detailed *in situ* analyses.

APPENDIX

LIST OF ACRONYMS²⁵

A_L	Lambert albedo
ADR	Ancillary Data Record
ADR TE	ADR Thermal Emission
ADR AC	ADR Atmospheric Correction
ADR CL	ADR Climatology
CAT	CRISM Analysis Tool
CRISM	Compact Reconnaissance Imaging Spectrometer for Mars (which is onboard MRO)
CRISM_LambertAlb	Name of the pipeline software system described here
DISORT	DIScrete-Ordinate-method Radiative Transfer
EDR	Experimental Data Record
EMI	Emission or Emergence angle

²⁵We also define some of these acronyms in the text, for ease of reading. They are defined here, in order to have a comprehensive list, for easy reference.

GRS	Gamma Ray Spectrometer (which is onboard Mars Odyssey)
HiRISE	High Resolution Imaging Science Experiment (which is onboard MRO)
HRSC	High-Resolution Stereo Camera (which is on-board Mars Express)
INC	Incidence angle
I/F	measured radiance divided by solar irradiance
L detector	Long-wavelength detector for CRISM
LUT	Lookup table
MARSIS	Mars Advanced Radar for Subsurface and Ionosphere Sounding (which is on-board Mars Express)
MGS	Mars Global Surveyor
MGS-TES ²⁶	Mars Global Surveyor—Thermal Emission Spectrometer
MOC	Mars Orbiter Camera (which was onboard MGS)
MOLA	Mars Orbiter Laser Altimeter (which was onboard Mars Global Surveyor)
MRDR	Mapped Reduced Data Record
MRO	Mars Reconnaissance Orbiter
MSL	Mars Science Laboratory
MSP	Multispectral
OMEGA	Observatoire pour la Minéralogie, l'Eau, les Glaces et l'Activité (which is on-board Mars Express)
PHI	Azimuthal angle
PDS	Planetary Data System
PFS	Planetary Fourier Spectrometer (which is onboard Mars Express)
S detector	Short-wavelength detector for CRISM
SHARAD	Shallow Radar sounder (which is onboard MRO)
THEMIS	Thermal Emission Imaging System (which was onboard Mars Odyssey)
TRDR	Targeted Reduced Data Record
USGS	U.S. Geological Survey
VL	Viking Lander

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²⁶This acronym is used instead of "TES" in order to distinguish this instrument from instruments of the same acronym on other spacecraft (i.e., the Tropospheric Emission Spectrometer, on-board the Aura satellite, orbiting the Earth).

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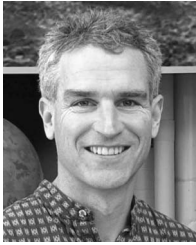


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